

Solar Neutrinos and the Decaying Neutrino Hypothesis

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We explore, mostly using data from solar neutrino experiments, the hypothesis that the neutrino mass eigenstates are unstable. We find that, by combining ^8B solar neutrino data with those on ^7Be and lower-energy solar neutrinos, one obtains a mostly model-independent bound on both the ν_1 and ν_2 lifetimes. We comment on whether a nonzero neutrino decay width can improve the compatibility of the solar neutrino data with the massive neutrino hypothesis.

The discovery of distinct nonzero neutrino masses and nontrivial lepton mixing opened the door to several fundamental questions that revolve around the properties of the neutral leptons. Here we concentrate on what, experimentally and model-independently, is known about the neutrino lifetime.

In the absence of interactions and degrees of freedom beyond those of the Standard Model, the two heaviest neutrinos – ν_2 and ν_3 (ν_1 and ν_2) in the case of the so-called normal (inverted) neutrino mass hierarchy [1] – are unstable, decaying into lighter neutrinos and photons ($\nu_i \rightarrow \nu_j \nu_k \nu_l$ or $\nu_i \rightarrow \nu_j + \gamma$, where $i, j, k, l = 1, 2, 3$). The associated lifetimes, given the tiny neutrino masses, are longer than 10^{37} years – much longer than the age of the universe. The presence of new interactions, degrees of freedom, etc., can, of course, change the picture dramatically.

Experimental bounds on the lifetimes of the neutrinos are much shorter than those expected from the Standard Model minimally augmented to include nonzero neutrino masses. Consulting ‘The Review of Particle Physics’ [1], one encounters different bounds that span almost twenty orders of magnitude. Bounds on the neutrino magnetic moment, for example, translate into bounds on radiative neutrino decays [2]. Results from cosmic surveys sensitive to the expansion rate of the universe at different epochs are consistent with the existence of around three independent neutrino states, naively indicating that these do not decay within the span of billions of years. In order to translate measurements of the expansion rate of the universe into a bound on the neutrino lifetime, however, one must consider the nature of the neutrino decay process, since the daughters of the putative decay also contribute to the expansion rate of the universe and could, in principle, mimic the contributions of their parents [3, 4]. The observation of the effects of nonzero neutrino masses in cosmic surveys might change the picture significantly [5].

Model-independent bounds exist from experiments where the number of neutrinos produced in the source can be compared to the number of neutrinos detected some distance away. These include all neutrino oscillation experiments. Given a baseline L and a “beam” energy E , one expects to be sensitive to a neutrino decay

width Γ_i for ν_i with mass m_i such that

$$\Gamma_i m_i \frac{L}{E} \equiv d_i \frac{L}{E} = 5.07 \left(\frac{d_i}{\text{eV}^2} \right) \left(\frac{L}{\text{km}} \right) \left(\frac{\text{GeV}}{E} \right) \gtrsim 1. \quad (1)$$

Here, for convenience, we define $d_i \equiv \Gamma_i m_i$, which has dimensions of energy-squared, for two reasons. On one hand, all bounds discussed here are sensitive to d_i : it is not possible to disentangle the neutrino mass from its decay width, both being unknown. On the other hand, d_i , measured in eV^2 , can be easily and directly compared to the neutrino mass-squared differences that are measured in neutrino oscillation experiments and “compete” with the decay effects. For conversion purposes, $d_i = 10^{-11} \text{ eV}^2$ translates into a lifetime $\tau_i = 70 \mu\text{s}$ for a neutrino with mass $m_i = 1 \text{ eV}$.

Using Eq. (1), it is easy to naively estimate that long-baseline accelerator experiments like MINOS, T2K, and NuMI, with $L/E \sim 10^3 \text{ km/GeV}$, are sensitive to $d_i \gtrsim 10^{-4} \text{ eV}^2$, atmospheric neutrino experiments like SuperKamiokande, with $L/E \lesssim 10^5 \text{ km/GeV}$, are sensitive to $d_i \gtrsim 10^{-6} \text{ eV}^2$, and the KamLAND reactor neutrino experiment, with $L/E \lesssim 2 \times 10^4 \text{ km/GeV}$, is sensitive to $d_i \gtrsim 10^{-5} \text{ eV}^2$. Detailed analyses of atmospheric and MINOS data, for example, translate into $d_3 \lesssim 10^{-5} \text{ eV}^2$ [6] and $d_3 < 1.2 \times 10^{-4} \text{ eV}^2$ [7], respectively, assuming $d_1, d_2 \ll d_3$.

Astrophysical neutrinos, when directly observed in Earth-bound detectors, provide significantly more stringent bounds on some of the d_i . The observation of neutrinos from Supernova 1987A implies that at least one of the neutrino mass eigenstates made it from the explosion to the Earth and can be translated into $d_i < 1.2 \times 10^{-21} \text{ eV}^2$ for at least one $i = 1, 2, 3$ [8]. A very strong bound on at least one of the d_i can also be derived [9–14] from the current and future observations of ultra-high-energy neutrinos using the IceCube detector [15].

Solar neutrinos have $L/E \sim 10^{11} \text{ km/GeV}$ and hence are sensitive to $d_i \gtrsim 10^{-12} \text{ eV}^2$. The authors of [16] were the first to point out that the ^8B solar neutrino data translate into a very robust bound on $d_2 \lesssim 10^{-11} \text{ eV}^2$, mostly independent from d_1 and d_3 . In this letter, we revisit the impact of decaying neutrinos on solar neutrino data. Since the publication of [16], our understanding of solar neutrinos and neutrino properties improved significantly. More and more precise KamLAND data

not only confirmed the neutrino oscillation interpretation of solar neutrino data, but also provided a precision measurement of the “solar” mass-squared difference, $\Delta m_{12}^2 \equiv m_2^2 - m_1^2$, and a good independent measurement of the “solar” mixing angle θ_{12} [17]. Borexino data allow a precision measurement of ${}^7\text{Be}$ solar neutrinos, and a clean measurement of the pp solar neutrinos [18]. Finally, recent reactor [19–21] and accelerator data [22] have measured the “reactor angle” θ_{13} , revealing that it is nonzero but quite small, $\sin^2 \theta_{13} \sim 0.02$. We will argue that all this information allows one to place, almost model-independently, bounds on both d_1 and d_2 from solar neutrino data. These results, when combined with results from atmospheric neutrinos, allow one to unambiguously place bounds on all three d_i , $i = 1, 2, 3$, which are robust, mostly model independent, and do not depend on the values of the neutrino masses or the neutrino mass hierarchy.

We will show, *a posteriori*, that decay effects are negligible for the L/E values probed by the KamLAND experiment. This implies that the oscillation results obtained from KamLAND apply even if the neutrinos have a finite lifetime, including the fact that $\Delta m_{12}^2 \sim 10^{-4} \text{ eV}^2$ and $\sin^2 2\theta_{12} \sim 0.8$. This in turn implies that neutrino oscillations from the core to the edge of the Sun, to a very good approximation, satisfy the adiabatic approximation. Ignoring the (very small) day–night effect but taking into account that the different neutrinos can decay into final states not accessible to the different solar neutrino detectors, the probability $P_{e\alpha}$ that a neutrino with energy E born in the Sun as a ν_e is detected as a ν_α , $\alpha = e, \mu, \tau$ one astronomical unit L_\odot away from the Sun is

$$P_{e\alpha}(E) \simeq \sum_{i=1,2,3} p_{ei}(E) |U_{\alpha i}|^2 e^{-d_i L_\odot/E}, \quad (2)$$

where $U_{\alpha i}$, $i = 1, 2, 3$, are the elements of the neutrino mixing matrix, while $p_{ei}(E)$ are the probabilities that the neutrino exits the Sun as ν_i neutrino mass eigenstates. Strictly speaking, Eq. (2) is a good approximation when $d_i R_\odot/E \ll 1$ where R_\odot is the average solar radius. We will show that this is indeed the case for d_1 and d_2 , and we argue in the next paragraph that solar data are not sensitive to d_3 effects.

Given that $|\Delta m_{13}^2| \sim 2 \times 10^{-3} \text{ eV}^2$ – even if one includes nonzero neutrino decay widths [6, 7] – $p_{e3}(E) \simeq |U_{e3}|^2$ for all relevant solar neutrino energies, $E \in [100 \text{ keV}, 20 \text{ MeV}]$. Since $|U_{e3}|^2 = \sin^2 \theta_{13} \simeq 0.02$ is small, given the precision of the solar neutrino data, d_3 related effects are irrelevant. In other words, the solar data are consistent with all d_3 values. We anticipate that $d_3 \neq 0$ effects impact only very modestly the constraints on the other oscillation and decay parameters. Henceforth, we ignore θ_{13} effects – we formally set it to zero – and treat solar neutrino oscillations as if there were only two neutrinos, ν_e and ν_a (a for active).

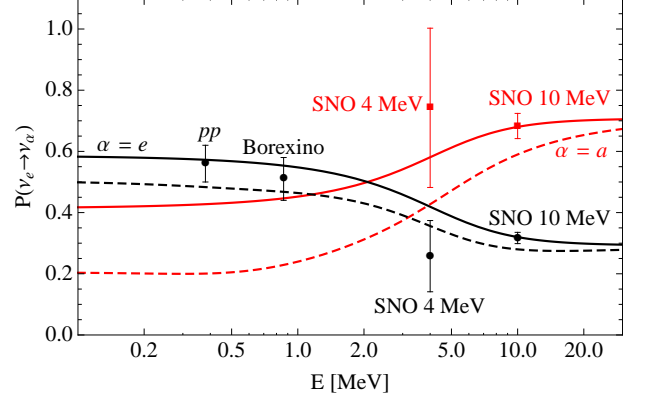


FIG. 1: P_{ee} (black) and P_{ea} (red) as a function of the solar neutrino energy, for $\sin^2 \theta_{12} = 0.29$, $\Delta m_{12}^2 = 7.5 \times 10^{-5} \text{ eV}^2$, $d_1 = 0$ and $d_2 = 0$ (solid) or $d_2 = 2 \times 10^{-12} \text{ eV}^2$ (dashed). Also depicted are the data points used to estimate the allowed values of d_1 and d_2 , including one sigma error bars. See text for details.

At high solar neutrino energies, $E \gtrsim 5 \text{ MeV}$, $p_{e2} \sim 1$, $p_{e1} \sim 0$, such that $P_{ee} \sim \sin^2 \theta_{12} e^{-d_2 L_\odot/E}$ and $P_{ea} \sim \cos^2 \theta_{12} e^{-d_2 L_\odot/E}$. ${}^8\text{B}$ solar neutrino data are hence very sensitive to d_2 but have little sensitivity to d_1 [16]. The most recent solar data from SNO [23] indicate a P_{ee} that decreases slowly as the neutrino energy decreases (as opposed to increasing, as predicted by the standard scenario, $d_1 = d_2 = 0$), a fact that is consistent with a judicious choice of d_2 . For illustrative purposes, Fig. 1 depicts P_{ee} and P_{ea} as a function of E , for $\sin^2 \theta_{12} = 0.29$, $\Delta m_{12}^2 = 7.5 \times 10^{-5} \text{ eV}^2$, $d_1 = 0$, and $d_2 = 0$ or $d_2 = 2 \times 10^{-12} \text{ eV}^2$.

At low solar neutrino energies, $E \lesssim 1 \text{ MeV}$, solar neutrino oscillations are well approximated by simple, averaged-out vacuum oscillations such that $p_{e1} \sim \cos^2 \theta_{12}$, $p_{e2} \sim \sin^2 \theta_{12}$, and $P_{ee} \sim \cos^4 \theta_{12} e^{-d_1 L_\odot/E} + \sin^4 \theta_{12} e^{-d_2 L_\odot/E}$ and $P_{ea} \sim \sin^2 \theta_{12} \cos^2 \theta_{12} (e^{-d_2 L_\odot/E} + e^{-d_1 L_\odot/E})$. ${}^7\text{Be}$ and pp solar neutrino measurements are hence sensitive to both d_1 and d_2 . In isolation, the low-energy solar neutrino data can be used to place a bound on either d_1 or d_2 , but not both. This is easy to see: the data are consistent with $e^{-d_1 L_\odot/E} \rightarrow 0$ or $e^{-d_2 L_\odot/E} \rightarrow 0$ as long as one judiciously chooses $\sin^2 \theta_{12}$. For example, in the limit, say, $e^{-d_1 L_\odot/E} \rightarrow 0$, $P_{ee} \sim \sin^4 \theta_{12}$ and $P_{ea} \simeq \sin^2 \theta_{12} \cos^2 \theta_{12}$ can be made to fit the data, roughly, $P_{ee} \sim 0.55$ [18], by choosing $\sin^2 \theta_{12} = 0.75$ (in the “dark side” [24]), which is consistent with data from KamLAND [17]. This possibility, however, is ruled out by ${}^8\text{B}$ data, which “require” $\sin^2 \theta_{12} \sim 0.3$.

In summary, combined low and high energy solar neutrino data allow one to place nontrivial bounds on both d_1 and d_2 , i.e., the possibility that either $e^{-d_1 L_\odot/E} \rightarrow 0$ or $e^{-d_2 L_\odot/E} \rightarrow 0$ is ruled out. d_2 is mostly constrained by the ${}^8\text{B}$ data, while d_1 is mostly constrained by the ${}^7\text{Be}$

and pp data. Given the order-of-magnitude difference between the neutrino energies, we anticipate the d_1 bound to be, roughly, an order of magnitude stronger than the d_2 bound.

In order to estimate the upper bounds on d_1 and d_2 , we perform a simple χ^2 fit to $\sin^2 \theta_{12}, d_1, d_2$, fixing $\Delta m_{12}^2 = 7.5 \times 10^{-5} \text{ eV}^2$, the best fit from KamLAND, and setting $\sin^2 \theta_{13} = 0$, as discussed earlier. Given that KamLAND provides the dominant contribution to the measurement of Δm_{12}^2 , this is a very reasonable approximation. Since in the decaying-neutrinos scenario $P_{ee} + P_{ea} \leq 1$, we need to consider separately the information on the electron and the “active” neutrino components of the solar neutrino flux. In detail, we include the following experimental information, depicted in Fig. 1:

- $P_{ee} = 0.56 \pm 0.06$ for $E = 380 \text{ keV}$, as extracted from a combined fit to Borexino and low-energy neutrino data [18]. This analysis is performed, effectively, by using Borexino data in order to establish the oscillated ${}^7\text{Be}$ neutrino flux and hence extract the pp neutrino flux from other data. This procedure depends only weakly on the hypothesis that $P_{ee} + P_{ea} = 1$. This result is consistent with Borexino’s recent independent measurement of the low energy solar neutrino flux [25].
- $s = P_{ee} + rP_{ea} = 0.62 \pm 0.05$ for $E = 862 \text{ keV}$ from the Borexino data [18]. $r = 0.22$ is the ratio of the $\nu_e + e$ to the $\nu_a + e$ elastic scattering cross-sections at ${}^7\text{Be}$ neutrino energies.
- SNO performed a detailed measurement of P_{ee} as a function of energy [23]. We choose P_{ee} values at $E = 4 \text{ MeV}$ and $E = 10 \text{ MeV}$, $P_{ee} = 0.26 \pm 0.12$ and $P_{ee} = 0.32 \pm 0.02$, respectively, as representatives of the SNO data. These points are chosen in order to both capture the statistical power of the SNO experiment and to include some of the shape information. A proper treatment of the SNO data, including all different observables, correlations, etc., can only be handled by the Collaboration itself. We verify that, in the case $d_1 = d_2 = 0$, our extracted best fit value for $\sin^2 \theta_{12}$ and the associated one sigma error bar are in good agreement with the most recent global analyses of neutrino data [26].
- The SNO experiment is also sensitive to the presence of a ν_a flux from the Sun thanks to its neutral current and $\nu + e$ elastic scattering measurements. It is, therefore, possible to measure P_{ee} and P_{ea} as a function of energy with SNO data (see, for example, [28]). Ref. [23], however, does not discuss the independent extraction of P_{ea} from the data, replacing it instead by $1 - P_{ee}$. Here we estimate the extracted value of P_{ea} from SNO data as follows. We define the central value using $P_{ea} = 1 - P_{ee}$

while fixing the one-sigma error bar on P_{ea} as that on P_{ee} , multiplied by $\sqrt{5}$. The factor of 5 is very close to the ratio of the elastic $\nu_e + e$ cross-section to that for $\nu_a + e$ at ${}^8\text{B}$ neutrino energies and agrees with the relative uncertainties for the electron and active neutrino fluxes measured by SNO in [28].

We note that SuperKamiokande also measures the neutrino flux using elastic neutrino–electron scattering (see, e.g., [27]). We do not include data from SuperKamiokande in our simplified fit as they mostly contribute to the measurement of P_{ea} – which we can only estimate here – and have a higher energy threshold than SNO data.

Fig. 2 depicts the result of our fit in the $d_1 \times d_2$ -plane, obtained after marginalizing over $\sin^2 \theta_{12}$. The best fit point is $d_1 = 3.4 \times 10^{-19} \text{ eV}^2, d_2 = 1.6 \times 10^{-13} \text{ eV}^2$ and the hypothesis $d_1 = d_2 = 0$ fits the data quite well. At the two-sigma confidence level, $d_1 < 1.6 \times 10^{-13} \text{ eV}^2$ and $d_2 < 9.3 \times 10^{-13} \text{ eV}^2$, in agreement with the naive estimates discussed above. The constraints above justify the approximations that led to Eq. (2), especially $d_{1,2}R_\odot/E \ll 1$ for all solar neutrino energies. Our result indicates that the neutrino decay hypothesis does not allow for a fit to the solar data that is significantly better than the standard “large mixing angle solution,” mostly due to the low energy ${}^7\text{Be}$ and pp neutrino measurements. We emphasize, however, that a detailed analysis of all solar neutrino data including the neutrino decay hypothesis is best left to the experimental collaborations, and that a reanalysis of the SNO data – one that treats both $P_{ee}(E)$ and $P_{ea}(E)$ as independent functions – as a function of energy is required. We hope our results encourage the pursuit of such an analysis.

In summary, we have argued that the solar neutrino data, combined with reactor data, allow one to place mostly model-independent bounds on the lifetimes of ν_1 and ν_2 . Using a subset of the solar neutrino data and the data from KamLAND, we estimate that $d_1 < 1.6 \times 10^{-13} \text{ eV}^2$ and $d_2 < 9.3 \times 10^{-13} \text{ eV}^2$ at the two-sigma confidence level. A complete analysis would reveal exactly where these bounds lie. As a “by-product” of our analysis, the atmospheric neutrino bound discussed in [6] applies, robustly, to ν_3 : $d_3 \lesssim 10^{-5} \text{ eV}^2$. Along with solar and atmospheric neutrino data, the only other robust bound comes from SN1987A which translates, as discussed earlier, into $d_i < 1.2 \times 10^{-21}$ for one of the three mass-eigenstates, most likely ν_1 or ν_2 .

The bounds are ‘mostly model-independent’ in the following sense. They are independent from the values of the neutrino masses themselves, and apply for both mass hierarchies. No assumption is made regarding the nature of the neutrino – Majorana or Dirac – or of the daughter particles into which the neutrinos would be decaying. We are assuming, however, that if the decay-daughters were to consist of lighter active neutrinos, these would

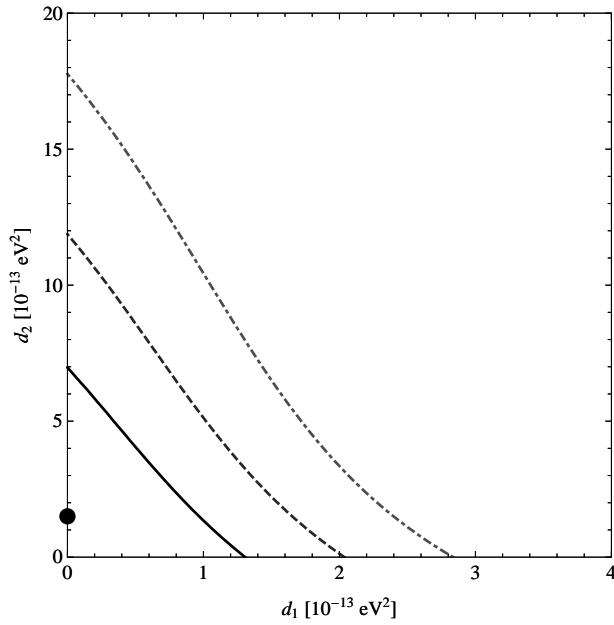


FIG. 2: Allowed values of d_1 and d_2 at the one, two and three sigma. The best fit point $d_1 = 3.4 \times 10^{-19} \text{ eV}^2$, $d_2 = 1.6 \times 10^{-13} \text{ eV}^2$ is indicated by a dot.

not leave a significant imprint in the detectors under consideration, i.e., they don't "look" like the parent neutrinos. This is a modest assumption. Daughter neutrinos from neutrino decay have, necessarily, less energy than their parents, and only those that decay along the flight-path of the parent make it to the detector. We also do not allow for the possibility, recently discussed in a more generic context [29], that the neutrino decay hypothesis translates into more mixing parameters, i.e., that the neutrino mass and decay eigenstates are not the same.

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- [2] C. Broggini, C. Giunti and A. Studenikin, *Adv. High Energy Phys.* **2012**, 459526 (2012) [arXiv:1207.3980 [hep-ph]].
- [3] J.F. Beacom, N.F. Bell and S. Dodelson, *Phys. Rev. Lett.* **93**, 121302 (2004) [astro-ph/0404585].
- [4] S. Hannestad and G. Raffelt, *Phys. Rev. D* **72**, 103514 (2005) [hep-ph/0509278].
- [5] P.D. Serpico, *Phys. Rev. Lett.* **98**, 171301 (2007) [astro-ph/0701699].
- [6] M.C. Gonzalez-Garcia and M. Maltoni, *Phys. Lett. B* **663**, 405 (2008) [arXiv:0802.3699 [hep-ph]].
- [7] R.A. Gomes, A.L.G. Gomes and O.L.G. Peres, arXiv:1407.5640 [hep-ph].
- [8] J.A. Frieman, H.E. Haber and K. Freese, *Phys. Lett. B* **200**, 115 (1988).
- [9] J.F. Beacom, N.F. Bell, D. Hooper, S. Pakvasa and T.J. Weiler, *Phys. Rev. Lett.* **90**, 181301 (2003) [hep-ph/0211305].
- [10] D. Meloni and T. Ohlsson, *Phys. Rev. D* **75**, 125017 (2007) [hep-ph/0612279].
- [11] P. Baerwald, M. Bustamante and W. Winter, *JCAP* **1210**, 020 (2012) [arXiv:1208.4600 [astro-ph.CO]].
- [12] S. Pakvasa, A. Joshipura and S. Mohanty, *Phys. Rev. Lett.* **110**, 171802 (2013) [arXiv:1209.5630 [hep-ph]].
- [13] L. Dorame, O.G. Miranda and J.W.F. Valle, arXiv:1303.4891 [hep-ph].
- [14] L. Fu and C.M. Ho, arXiv:1407.1090 [hep-ph].
- [15] M.G. Aartsen *et al.* [IceCube Collaboration], *Phys. Rev. Lett.* **111**, 021103 (2013) [arXiv:1304.5356 [astro-ph.HE]].
- [16] J.F. Beacom and N.F. Bell, *Phys. Rev. D* **65**, 113009 (2002) [hep-ph/0204111].
- [17] A. Gando *et al.* [KamLAND Collaboration], *Phys. Rev. D* **83**, 052002 (2011) [arXiv:1009.4771 [hep-ex]].
- [18] G. Bellini, J. Benziger, D. Bick, S. Bonetti, G. Bonfini, M. Buizza Avanzini, B. Caccianiga and L. Cadonati *et al.*, *Phys. Rev. Lett.* **107**, 141302 (2011) [arXiv:1104.1816 [hep-ex]].
- [19] F.P. An *et al.* [Daya Bay Collaboration], *Phys. Rev. Lett.* **112**, 061801 (2014) [arXiv:1310.6732 [hep-ex]].
- [20] J.K. Ahn *et al.* [RENO Collaboration], *Phys. Rev. Lett.* **108**, 191802 (2012) [arXiv:1204.0626 [hep-ex]].
- [21] Y. Abe *et al.* [Double Chooz Collaboration], *Phys. Rev. D* **86**, 052008 (2012) [arXiv:1207.6632 [hep-ex]].
- [22] K. Abe *et al.* [T2K Collaboration], *Phys. Rev. Lett.* **112**, 061802 (2014) [arXiv:1311.4750 [hep-ex]].
- [23] B. Aharmim *et al.* [SNO Collaboration], *Phys. Rev. C* **88**, 025501 (2013) [arXiv:1109.0763 [nucl-ex]].
- [24] A. de Gouvêa, A. Friedland and H. Murayama, *Phys. Lett. B* **490**, 125 (2000) [hep-ph/0002064].
- [25] G. Bellini *et al.* [BOREXINO Collaboration], *Nature* **512**, no. 7515, 383 (2014).
- [26] M.C. Gonzalez-Garcia, M. Maltoni and T. Schwetz, arXiv:1409.5439 [hep-ph].
- [27] K. Abe *et al.* [Super-Kamiokande Collaboration], *Phys. Rev. D* **83**, 052010 (2011) [arXiv:1010.0118 [hep-ex]].
- [28] Q.R. Ahmad *et al.* [SNO Collaboration], *Phys. Rev. Lett.* **89**, 011301 (2002) [nucl-ex/0204008].
- [29] J.M. Berryman, A. de Gouvêa, D. Hernández and R.L.N. Oliveira, arXiv:1407.6631 [hep-ph].

[1] K.A. Olive *et al.* [Particle Data Group Collaboration], *Chin. Phys. C* **38**, 090001 (2014).